The History of the Universe in 200 Words or Less

Quantum fluctuation. Inflation. Expansion. Strong nuclear interaction. Particle-antiparticle annihilation. Deuterium and helium production. Density perturbations. Recombination. Blackbody radiation. Local contraction. Cluster formation. Reionization? Violent relaxation. Virialization. Biased galaxy formation? Turbulent fragmentation. Contraction. Ionization. Compression. Opaque hydrogen. Massive star formation. Deuterium ignition. Hydrogen fusion. Hydrogen depletion. Core contraction. Envelope expansion. Helium fusion. Carbon, oxygen, and silicon fusion. Iron production. Implosion. Supernova explosion. Metals injection. Star formation. Supernova explosions. Star formation. Condensation. Planetesimal accretion. Planetary differentiation. Crust solidification. Volatile gas expulsion. Water condensation. Water dissociation. Ozone production. Ultraviolet absorption. Photosynthetic unicellular organisms. Oxidation. Mutation. Natural selection and evolution. Respiration. Cell differentiation. Sexual reproduction. Fossilization. Land exploration. Dinosaur extinction. Mammal expansion. Glaciation. Homo sapiens manifestation. Animal domestication. Food surplus production. Civilization! Innovation. Exploration. Religion. Warring nations. Empire creation and destruction. Exploration. Colonization. Taxation without representation. Revolution. Constitution. Election. Expansion. Industrialization. Rebellion. Emancipation Proclamation. Invention. Mass production. Urbanization. Immigration. World conflagration. League of Nations. Suffrage extension. Depression. World conflagration. Fission explosions. United Nations. Space exploration. Assassinations. Lunar excursions. Resignation. Computerization. World Trade Organization. Terrorism. Internet expansion. Reunification. Dissolution. World-Wide Web creation. Composition. Extrapolation?

ET: Astronomy 230

Section 1– MWF 1400-1450

134 Astronomy Building

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  T: 10:30-11:30 a.m.
  W: 3:00-4:30 p.m. or by appointment

This Class (Lecture 3):

Cosmology and the Origin of Elements

Next Class:

The Early Galaxy and the First Stars

HW1 due on Friday.

Outline

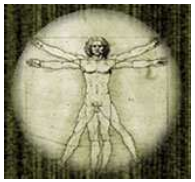
- Where did the atoms in our bodies come from?
- How old is the Universe?
- Big Bang Nucleosynthesis
- Cooling into normal stuff.
- The seeds of galaxies.
- What is the probable fate of the Universe?

The Universe

- Began with a Big Bang
  - 13.7 billion years ago
- Still expanding and cooling
  - The rate of expansion is known
- It is BIG
  - As far as we are concerned, it is infinite in any direction
- The universe is homogeneous and isotropic
  - **Homogeneous** - The same "stuff" everywhere
  - **Isotropic** - The same in all directions
- Our place in the Universe is not special
  - Extension of the Copernican revolution
- The center of the Universe is everywhere or nowhere!

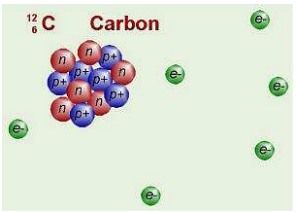
Defining Life

As we will discuss later, defining life is very difficult. Traditional attributes of life define it as:



Elements of Life

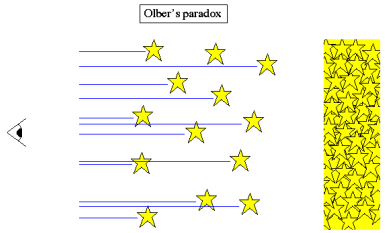
- Carbon is the most important element in life on Earth with oxygen and nitrogen coming in a close second. And there is a lot of hydrogen. But where did they come from?
- To understand this question, we need to address the origin of the Universe and the elements crucial to life.
- In other words, Cosmology.



http://biology.clc.uc.edu/courses/bio104/atom-h2o.htm

The Night Sky: **Olber's Paradox**

- Why is the night sky not bright with light?
- If the Universe is infinite, why don't we see light everywhere from all the stars.
- Even if dust blocked the light, it would heat up and emit in the optical too.
- The Universe has not existed forever. It must have started from something.



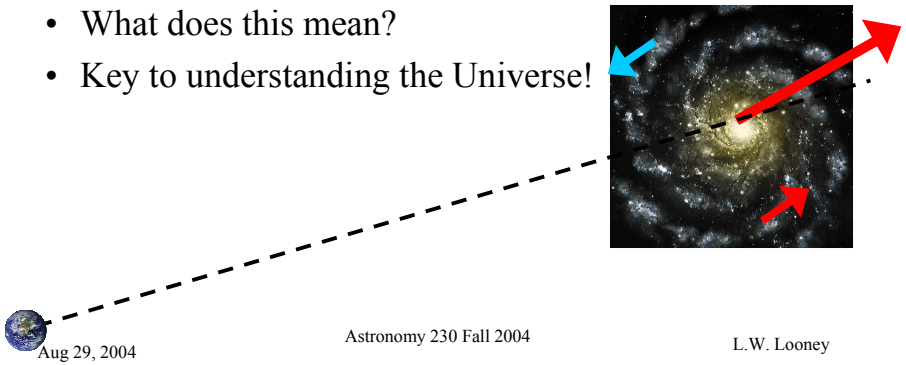
How are Galaxies Moving?

It's 1928 and Edwin Hubble is measuring how galaxies move. What does he find?

a) More galaxies receding than approaching.
b) More galaxies approaching than receding.
c) About equal numbers of each.

Redshift of Galaxies

- Most galaxies are moving away from us.
- The farther away, the faster they are moving away.
- Or $V = H_o \times D$
  - H_o = 72 km/s /Mpc 1 parsec = 3.262 lyr
- What does this mean?
- Key to understanding the Universe!



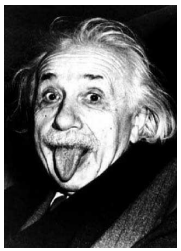
Apply it?

- In a homogenous Universe, what does the farther away the faster they move away mean?
- Draw it.

Interpretation: View of the Universe

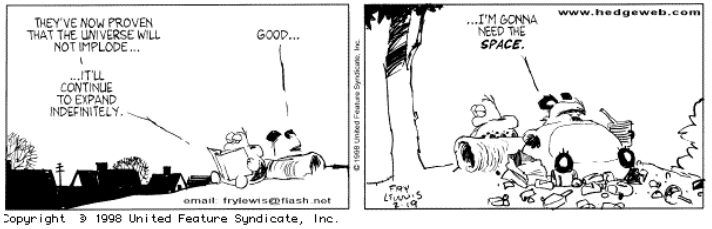
Egoist view– We are at the center of the Universe.

Einstein's view– The Universe is expanding, and there is no center!



The Expanding Universe

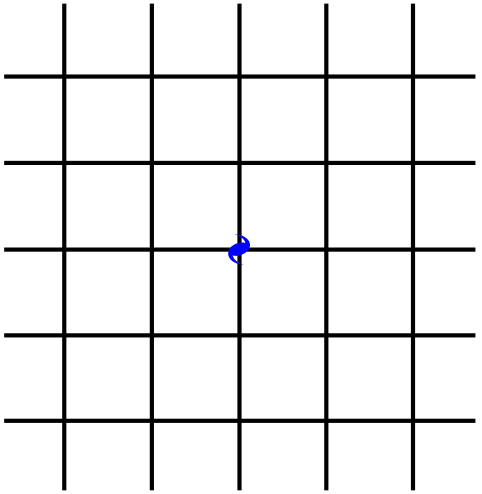
- To describe the motion of all the galaxies in the Universe, we use General Relativity (due to the gravity effects)
- General Relativity predicts that we live in an *expanding Universe*.
  - Einstein didn't buy it at first, so made a cosmological constant to get rid of it.
- In other words, space is stretching in all directions. This completely explains Hubble's Law.



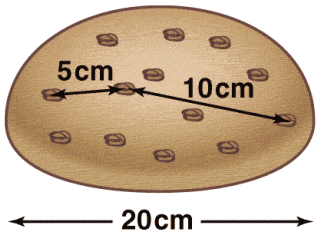
What do you think?

- The Universe is expanding, how do you feel about that?

Dude, The Universe is Expanding.



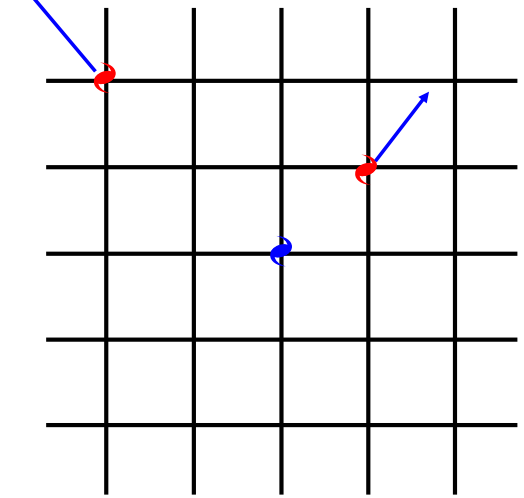
Analogy– Raisin Bread



MAP980404

Raisins stay the same size.

Wow. The Universe is Expanding.



~~Expanding into What?~~

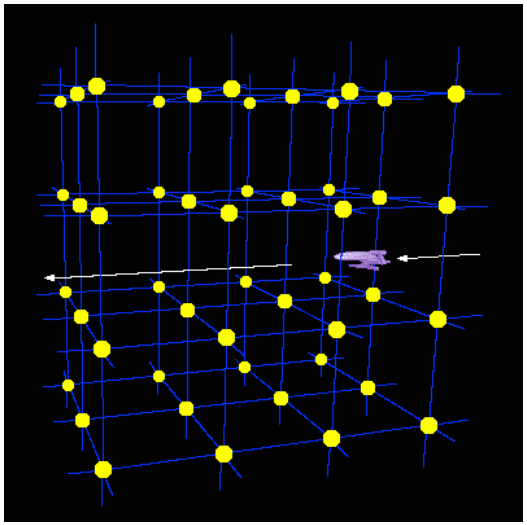
Reality

- The analogies are just to help us visualize, don't get stuck in the specifics.
- The Universe has no center.
- The Universe has no edge.
- Concept of time and space began with the Universe, can not apply the concepts so easily.

http://universe.gsfc.nasa.gov/images/reach-for-the-universe.jpg

The Edge of the Universe?

- If the Universe consisted of only 48 stars?
- The spaceship, would never really see the edge of the Universe.



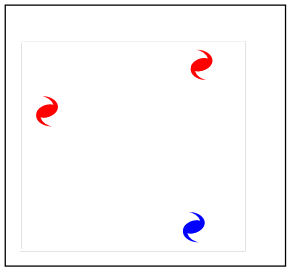
http://www.anzwers.org/free/universe/bigbang.html

Living in an Expanding Universe

Consider a large "box" containing many galaxies

- Total mass in box today: M_{today}
- Total volume in box today: V_{today}
- ***Density today*** $= M_{today} / V_{today}$

The Universe box



Tomorrow

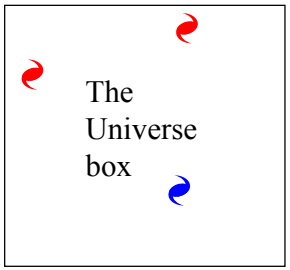
How does the density of the Universe change with time?

Living in an Expanding Universe

How does the density of the Universe change with time? As Universe expands:

- $M_{tomorrow}$ stays the same
- $V_{tomorrow}$ becomes larger
- Density $M_{tomorrow}/V_{tomorrow} \Rightarrow$ ***smaller***

$M_{tomorrow}/V_{tomorrow} < M_{today}/V_{today}$



Density changes with time!

- Universe was denser the past
- Universe will be less dense in future

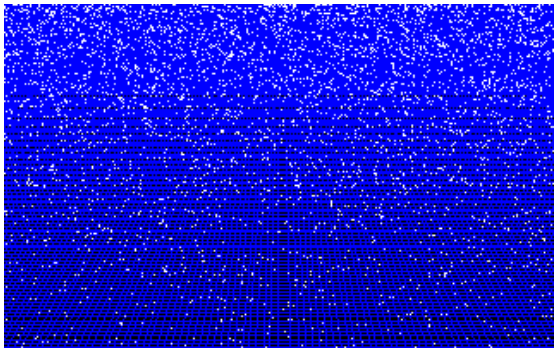
Putting it all together:

1. Earlier Universe was more dense
2. Earlier Universe was hotter.
3. The Universe is expanding.

The origin of the Universe can be described by the idea of the Big Bang. Where did the Big Bang happen? Remember the Universe is homogenous & isotropic.

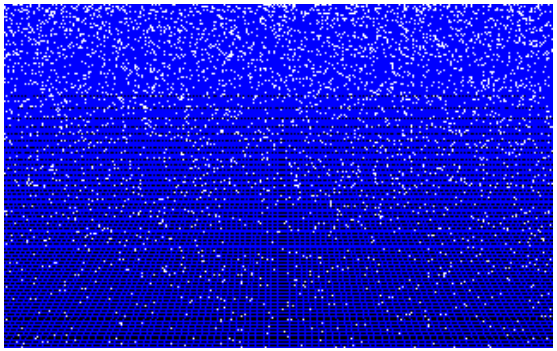
The Big Bang

- Occurred everywhere at once.
- No special points or locals
- Expansion of **all** space
- Not an explosion into empty space.



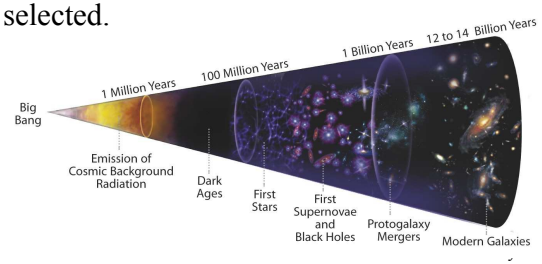
The Big Bang

- Big Bang has no center
- Happened everywhere
- Wherever you go, there was the big bang
- So as we talk about the very dense early universe, remember that we are talking about what happened not just far away at the edge of the Universe, but **right here!** ...smooshed up small, but still **right here!**

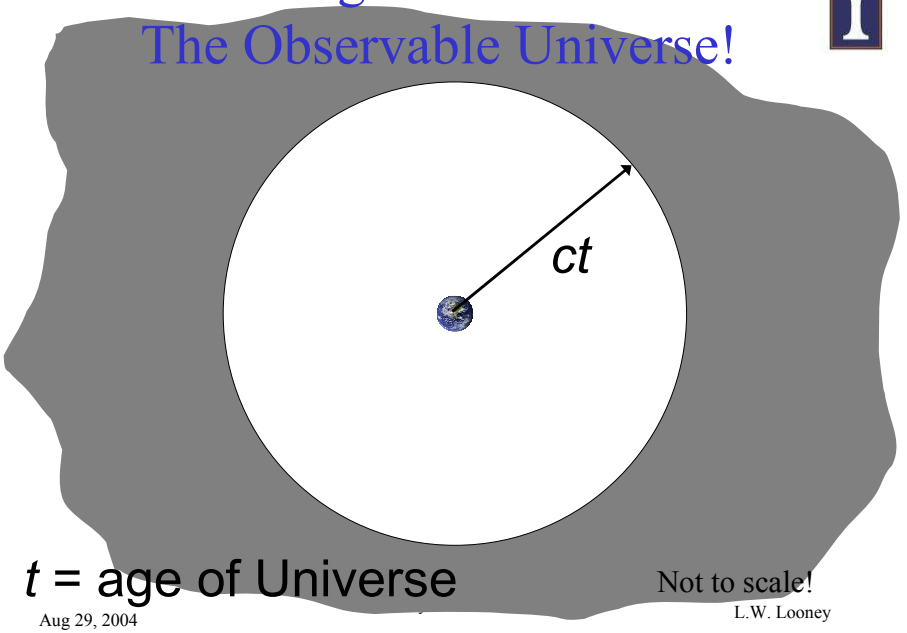


The Big Bang

- In the 1940s, extrapolating on Hubble's Law, George Gamow proposed the the universe began in a colossal "explosion" of expansion.
- In the 1950s, the term BIG BANG was coined by an unconvinced Sir Fred Hoyle who tried to ridicule it.
- In the 1990s, there was an international competition to rename the BIG BANG with a more appropriate name, but no new name was selected.



Looking Back in Time: The Observable Universe!

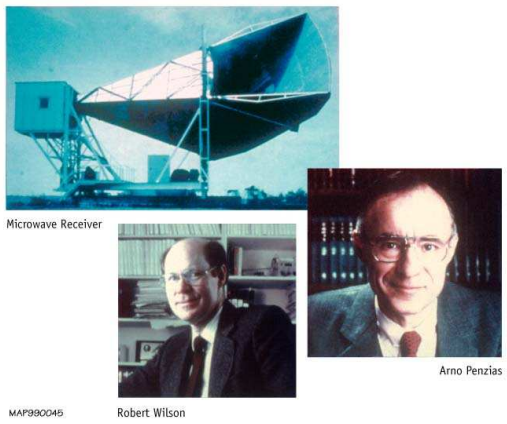


t = age of Universe

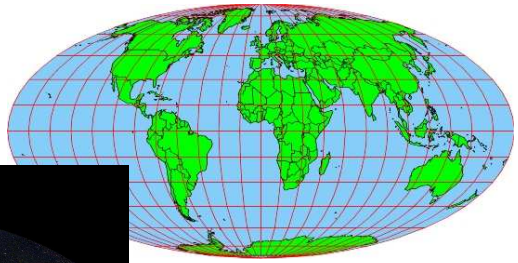
Not to scale!

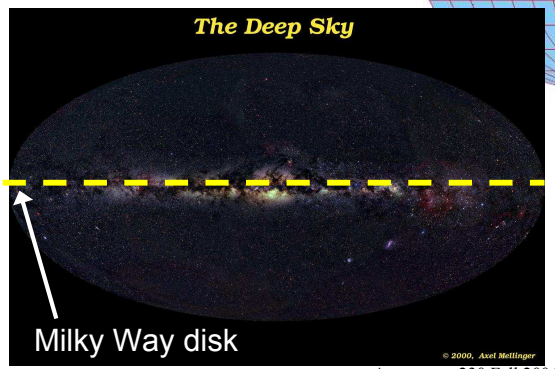
The Early Universe was ***HOT***!

- If the early Universe was so hot, we should be able to see it glowing. Right?
- Yep, we do! But, as the Universe expanded, it shifted colors down to the microwave.
- Now, it is called the Cosmic Microwave Background (CMB).
- First detected by Robert Wilson and Arno Penzias.

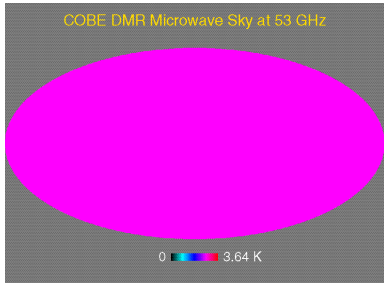


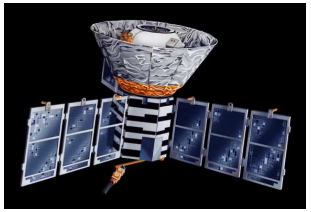
How to Understand Sky Maps





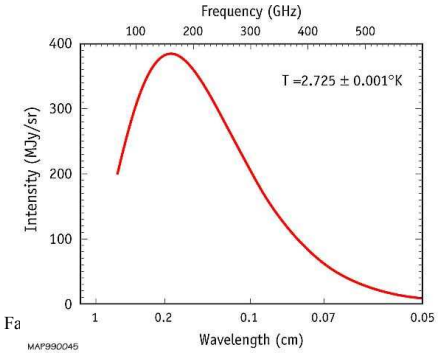
A Rather Uniform Blackbody

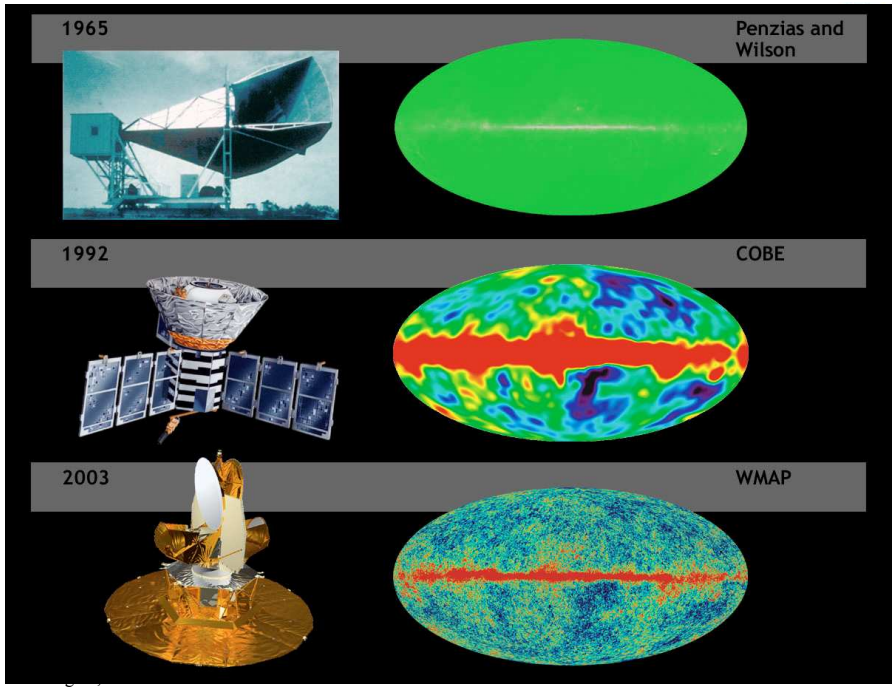




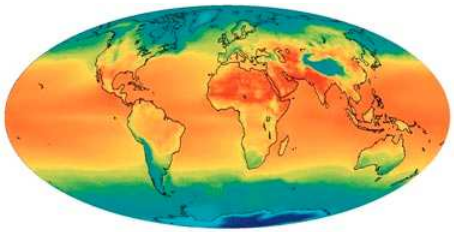
Cosmic Background Explorer (COBE) satellite (launched 1989)

$$T \approx 3\ \mathrm{K}$$



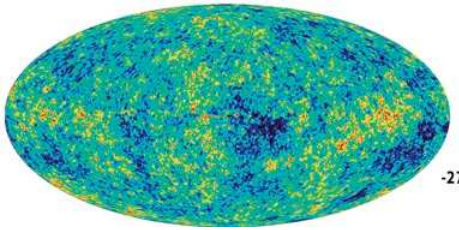


WMAP took a "baby picture" of the Universe– only 400000 yrs old.



Earth Temperatures

-63° -13° 37°
Centigrade
June 1992



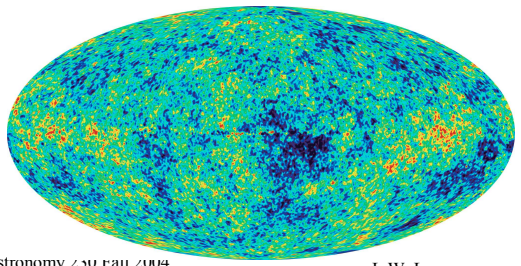
Microwave Sky Temperatures

-270.4252° -270.4250° -270.4248°
Centigrade
380,000 Years after Big Bang

The Seeds of Galaxies

These small perturbations in temperature are the fluctuations (smaller than 1 in a 1000) that caused the large scale structures we see today. This is what formed galaxies. All of this happened only 400,000 years after the Big Bang.

http://map.gsfc.nasa.gov/m_ig/030651/030651b.mov

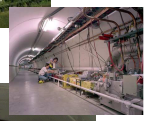


THE VERY EARLY UNIVERSE

Since Big Bang works well so far, we have confidence to think about times earlier still:

$t << 1$ sec !

- Temperature and energies are ***ultrahigh***



Q: How to probe such high energies?
Hint: it's in the Great State of Illinois

Fermilab

INNER SPACE / OUTER SPACE

Fermilab is a telescope!

Probes conditions in Universe at 10^{-12} s

Universe was 10^{12} K hot!

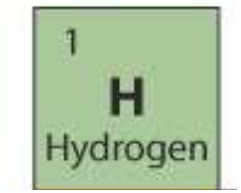
…but also…

"The Universe is the poor man's accelerator"

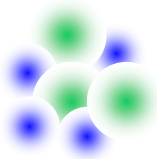
Probes conditions inaccessible at laboratories

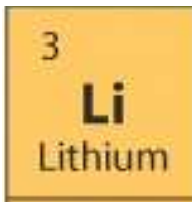
Big Bang Nucleosynthesis

When the Universe was 3 seconds old, the temperature fell to 10^9 K and protons and neutrons can "shack-up" to form the first light elements.

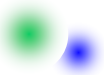


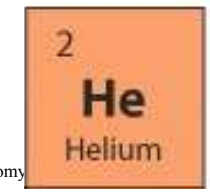
(a proton)

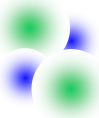




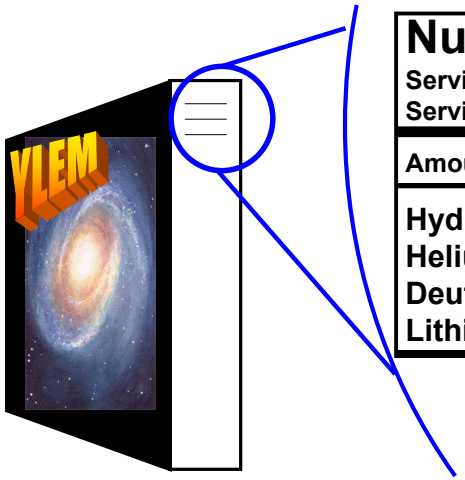
Also: Deuterium







End Result: Big Bang Correctly Predicts Abundances



Nutrition Facts

Serving Size 1 g
Servings Per Universe many many

Amount Per Serving

Hydrogen	**0.75 g**
Helium	**0.25 g**
Deuterium	**10^{-4} g**
Lithium, etc.	**10^{-10} g**

A Brief History of Time

Aug 29, 2004

Astronomy 230 Fall 2004

L.W. Looney